



# Updates from Simons Observatory (SO)

Remington G. Gerras

University of Southern California

AAS 243, New Orleans, LA



# Overview

- Simons Observatory (SO) Introduction
- SO on-site updates
  - Large Aperture Telescope and Receiver (LAT/LATR)
  - Small Aperture Telescopes (SATs)
- SO pipeline updates
- SO Timeline and what's next

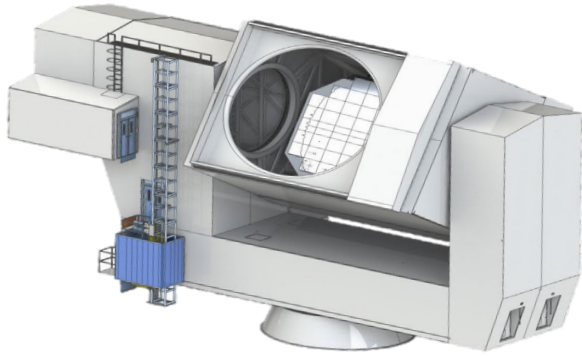
# Simons Observatory (SO)

- Atop Cerro Toco, Atacama Desert, Chile
  - 5200m above sea-level!
- Site as of 10/2023



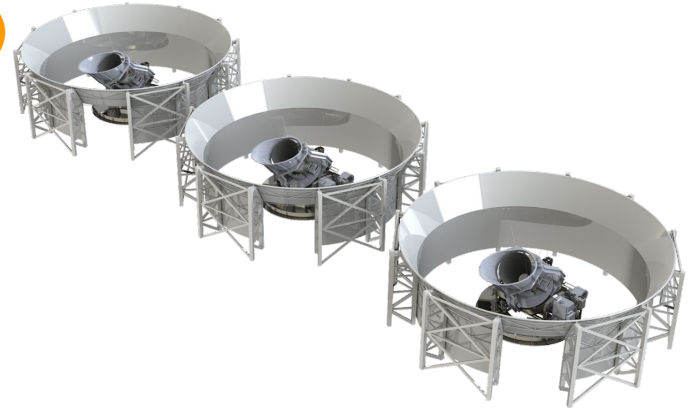
# SO Telescopes

## Large Aperture Telescope (LAT)



- 6m crossed-dragone telescope
- Large area scan strategy with high resolution ( $\sim$ arc minute)
- 6 frequency bands: (27, 39, 90, 150, 220, 280) GHz
- 30k TES detectors

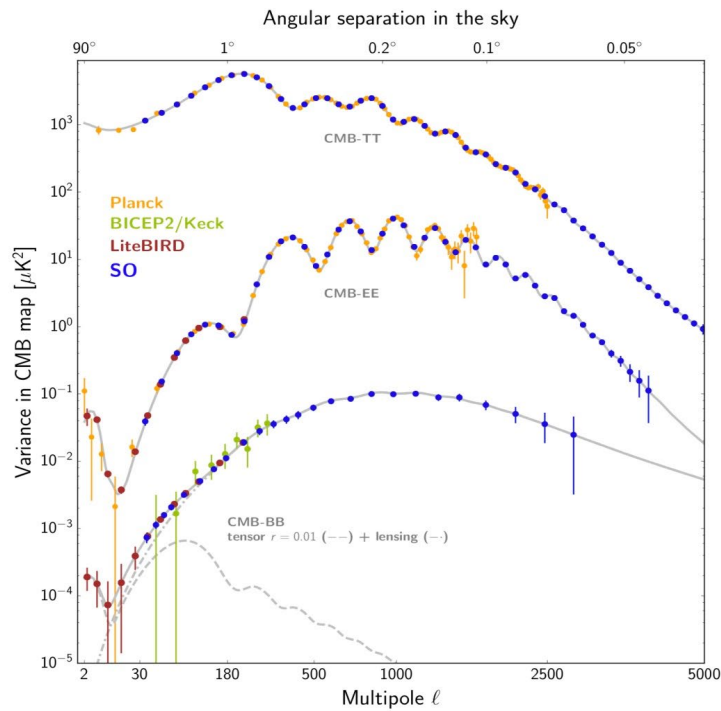
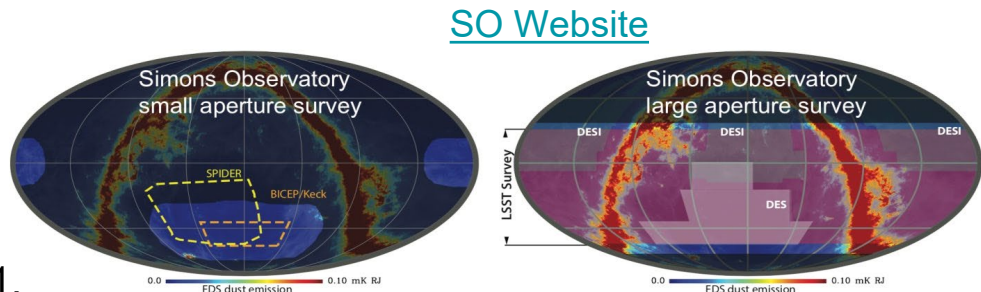
## 3x Small Aperture Telescopes (SATs)



- Refractive telescopes with 42cm aperture
- Deep scan with  $\sim$ degree resolution
- 6 frequency bands, same as LAT
  - Split amongst 3 SATs
- 30k TES detectors total

# Overview of SO

- Studies anisotropies in CMB
  - Science Goals:
    - Primordial B-modes, with tensor-to-scalar ratio,  $r < 0.01$ ,  $\sigma(r) < 0.003$
    - Neutrino Mass Hierarchy,  $N_{\text{eff}}$ , + many others! ([SO forecast paper](#))
- Nominally 4 telescopes sensitive to 6 frequency bands to observe extremely faint signal and to remove sources of galactic noise
  - LAT focuses on high  $\ell$  science
  - 3x SATs focus on primordial B-modes (low  $\ell$ )



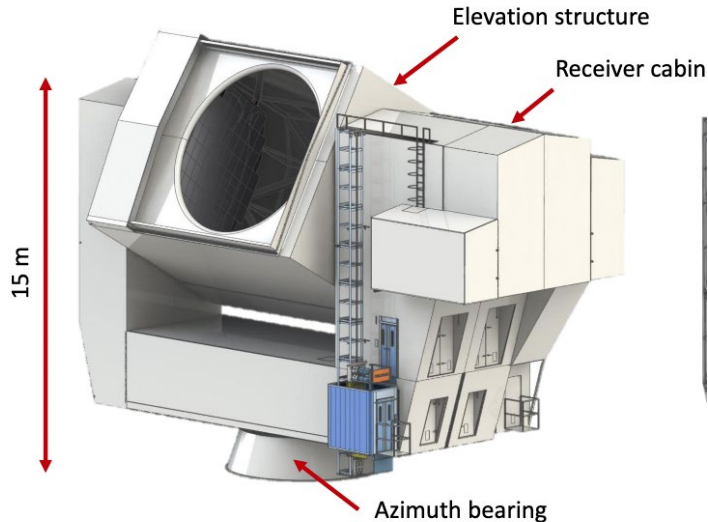
# Large Aperture Telescope (LAT)



Credit: Mark Devlin

# Large Aperture Telescope (LAT)

- Platform on site 06/2023
- Cryogenic Receiver installed 08/2023
  - Dark Tests performed last quarter of 2023
  - 2/7 optics tubes installed
- Mirrors to be installed midway through 2024
  - 7/7 optics tubes installed by 02/2024



Credit: Gakitzki SPIE 2018



# Small Aperture Telescope - Middle Frequency 1, 2 and Ultra High Frequency (SAT-MF1, MF2, UHF)



Credit: Sam Day-Weiss



Credit: Adrian Lee



# SAT-MF1

- Finished in-lab testing at UCSD
- Installed on Platform (08/2023)
- On-sky testing on-going



# SAT-MF2

- Finished in-lab testing at Princeton
- Installed on Platform (11/2023)
- Preparing for on-sky tests



# SAT-UHF

- Finished in-lab testing at Berkeley
- Shipped to Chile (12/2023)
- Begin unpacking and installation on platform
  - First quarter of 2024



# SO Pipeline Status

- SO makes use of multiple publicly available data processing pipelines
  - [so-lenspipe](#), [BBPower](#)
- Large-scale B-mode signal [pipeline validated using recent instrument data](#)

## The Simons Observatory: pipeline comparison and validation for large-scale *B*-modes

Kevin Wolz<sup>1,2 \*</sup>, Susanna Azzoni<sup>3,4</sup>, Carlos Hervías-Caimapo<sup>5,6</sup>, Josquin Errard<sup>7</sup>, Nicoletta Krachmalnicoff<sup>1,2,8</sup>, David Alonso<sup>3</sup>, Carlo Baccigalupi<sup>1,2,8</sup>, Antón Baleato Lizancos<sup>9,10</sup>, Michael L. Brown<sup>11</sup>, Erminia Calabrese<sup>12</sup>, Jens Chluba<sup>11</sup>, Jo Dunkley<sup>17,18</sup>, Giulio Fabbian<sup>12,13</sup>, Nicholas Galitzki<sup>14</sup>, Baptiste Jost<sup>7,15</sup>, Magdy Morshed<sup>7</sup>, and Federico Nati<sup>16</sup>

- LAT pipelines are being developed and tested on ACT data
  - [DR6 Gravitational Lensing Map and Cosmological Parameters](#)
  - [Mitigating the impact of extragalactic foregrounds for the DR6 CMB lensing analysis](#)

# Future of SO

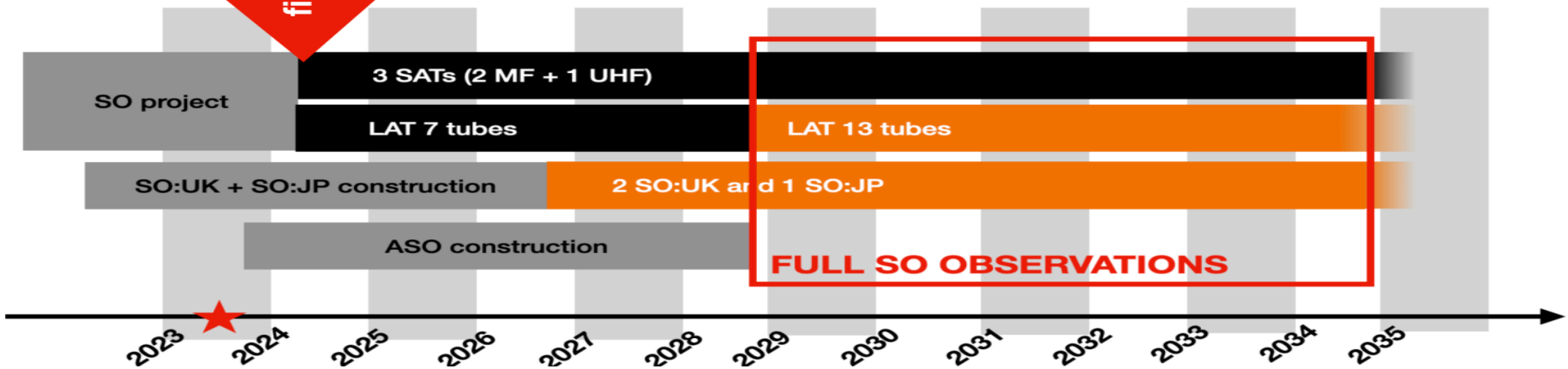
- SO:UK (United Kingdom), SO:JP (Japan)
  - 3 additional SATs (2x United Kingdom, 1x Japan)
  - 30k additional detectors
- Advanced SO (ASO)
  - Fully populate LAT Receiver with 6 additional optics tubes and 30k additional detectors
  - Additional 5 year observation time
  - Infrastructure upgrades to increase power efficiency via photovoltaic array



first light(s)

## SO timeline

Credit: Benjamin Beringue





End

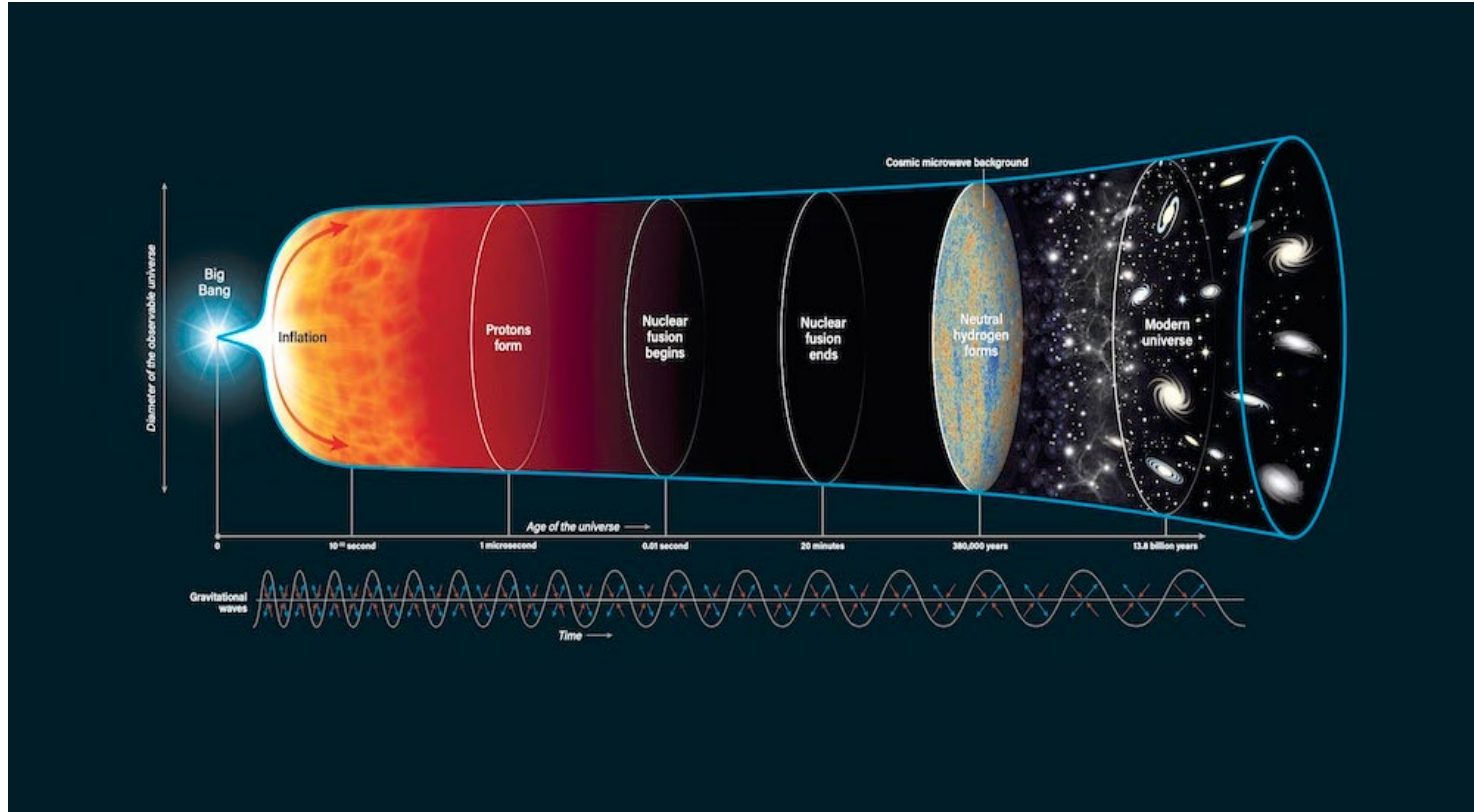
SDSC East Entrance

SDSC

# Extras

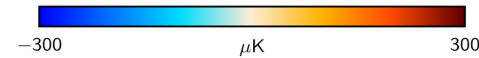
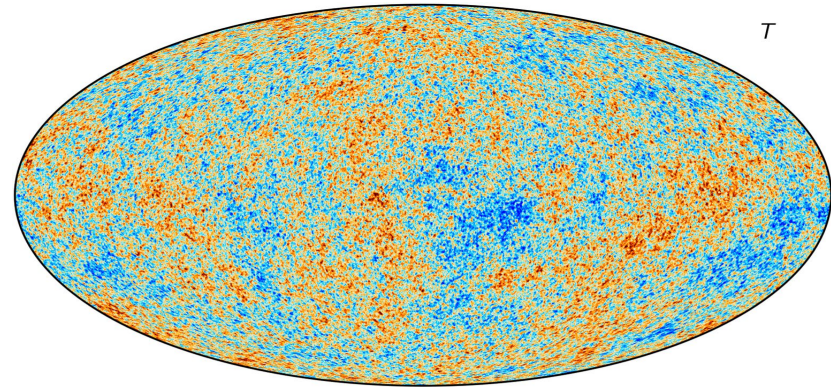
# (Quick) History of the Universe

- Density perturbations in very early universe are *inflated* to cosmic scales  
These seed cosmic structure we see today  
Density perturbations imprinted in the cosmic microwave background (CMB)

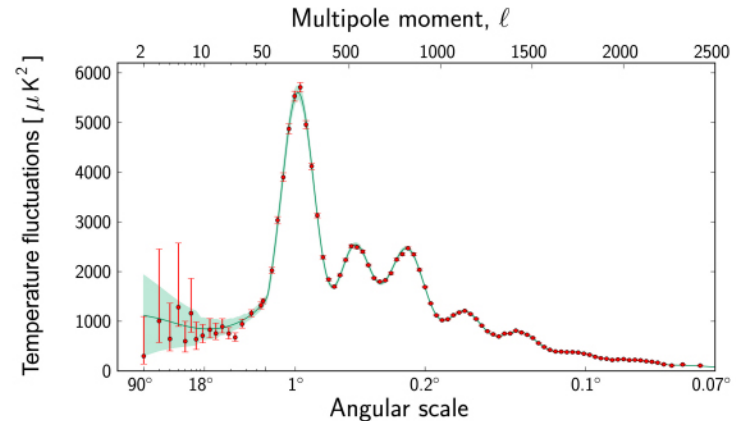


# Information from the CMB

- CMB is mostly uniform in temperature, but there are anisotropies. **Why?**
- Theory of inflation:
  - Temperature anisotropies caused by density perturbations
  - Polarization anisotropies as well!
- Study the CMB *statistics* via power spectra
  - In frequency and angular distance spaces



Credit: Planck Legacy Archive



Credit Planck Legacy Archive

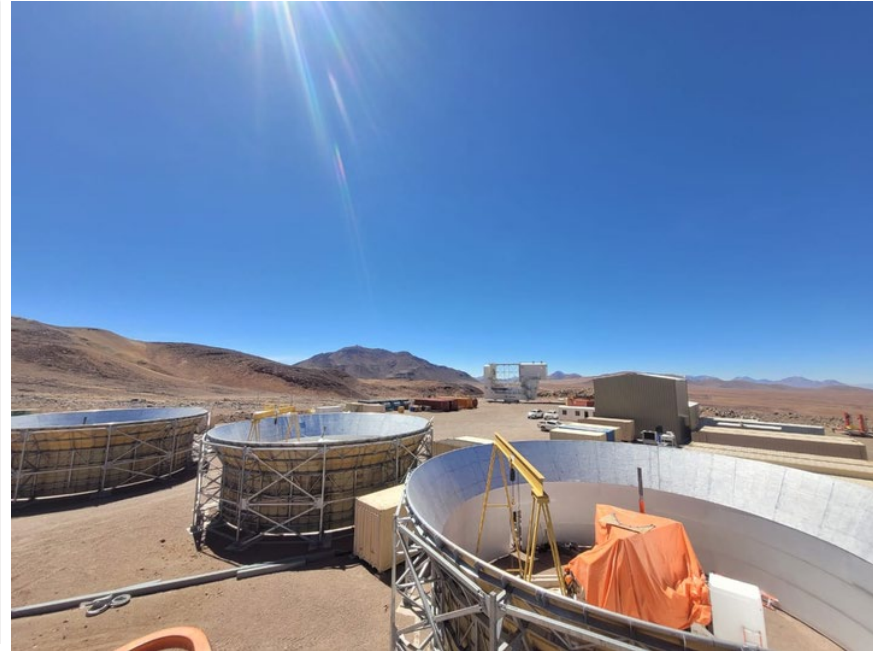
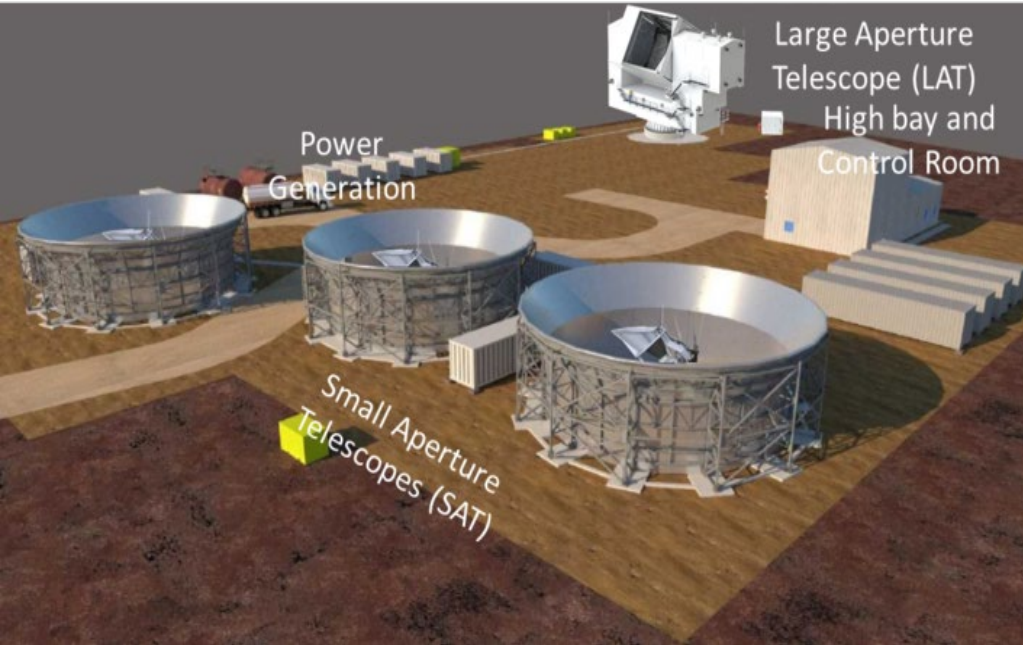


# SO-Nominal Key Science Goals

|  | Current <sup>b</sup> | SO-Nominal (2022-27) |                    | Method <sup>d</sup>                   |
|--|----------------------|----------------------|--------------------|---------------------------------------|
|  |                      | Baseline             | Goal               |                                       |
| <b>Primordial perturbations (§2.1)</b>       |                      |                      |                    |                                       |
| $r$ ( $A_L = 0.5$ )                          | 0.03                 | 0.003                | 0.002 <sup>e</sup> | BB + external delensing               |
| $n_s$  | 0.004                | 0.002                | 0.002              | TT/TE/EE                              |
| $e^{-2\tau} \mathcal{P}(k = 0.2/\text{Mpc})$ | 3%                   | 0.5%                 | 0.4%               | TT/TE/EE                              |
| $f_{\text{NL}}^{\text{local}}$               | 5                    | 3                    | 1                  | $\kappa\kappa \times \text{LSST-LSS}$ |
|  |                      | 2                    | 1                  | kSZ + LSST-LSS                        |
| <b>Relativistic species (§2.2)</b>           |                      |                      |                    |                                       |
| $N_{\text{eff}}$                             | 0.2                  | 0.07                 | 0.05               | TT/TE/EE + $\kappa\kappa$             |
| <b>Neutrino mass (§2.3)</b>                  |                      |                      |                    |                                       |
| $\Sigma m_\nu$ (eV, $\sigma(\tau) = 0.01$ )  | 0.1                  | 0.04                 | 0.03               | $\kappa\kappa$ + DESI-BAO             |
|  |                      | 0.04                 | 0.03               | tSZ-N $\times$ LSST-WL                |
| $\Sigma m_\nu$ (eV, $\sigma(\tau) = 0.002$ ) |                      | 0.03 <sup>f</sup>    | 0.02               | $\kappa\kappa$ + DESI-BAO + LB        |
|  |                      | 0.03                 | 0.02               | tSZ-N $\times$ LSST-WL + LB           |
| <b>Beyond standard model (§2.4)</b>          |                      |                      |                    |                                       |
| $\sigma_8(z = 1 - 2)$                        | 7%                   | 2%                   | 1%                 | $\kappa\kappa$ + LSST-LSS             |
|  |                      | 2%                   | 1%                 | tSZ-N $\times$ LSST-WL                |
| $H_0$ ( $\Lambda\text{CDM}$ )                | 0.5                  | 0.4                  | 0.3                | TT/TE/EE + $\kappa\kappa$             |
| <b>Galaxy evolution (§2.5)</b>               |                      |                      |                    |                                       |
| $\eta_{\text{feedback}}$                     | 50-100%              | 3%                   | 2%                 | kSZ + tSZ + DESI                      |
| $p_{\text{nt}}$                              | 50-100%              | 8%                   | 5%                 | kSZ + tSZ + DESI                      |
| <b>Reionization (§2.6)</b>                   |                      |                      |                    |                                       |
| $\Delta z$                                   | 1.4                  | 0.4                  | 0.3                | TT (kSZ)                              |

From Astro2020 talk (A. Lee, S. Staggs)

# Render of site to Current site status



From Mark Devlin; SO

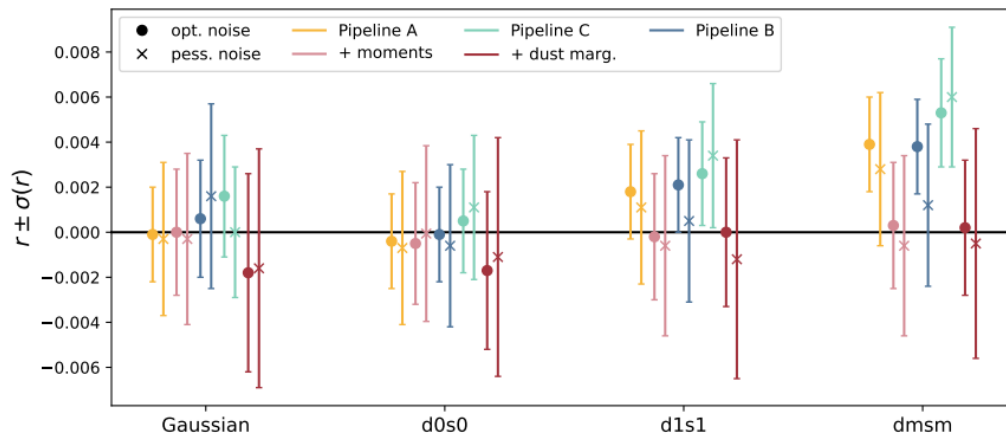
# Advanced Simons Observatory (ASO)

- Doubles mapping speed and increases number of detectors
  - Fully populate LAT Receiver (from 7/13 to 13/13 Optics Tubes, additional 30k detectors)
- Adds 5 years of observations

|  | SO-Nominal<br>(Goal)        | SO-Enhanced<br>(Goal)       |
|--|-----------------------------|-----------------------------|
| <b>Lensing and SZ (LAT)</b>  |                             |                             |
| Minimal neutrino mass detection ( $\Sigma m_\nu=0.06$ eV) <sup>a</sup> | $3\sigma$                   | $4\sigma$                   |
| Lensing detection (polarization-only)                                  | $160\sigma$ ( $110\sigma$ ) | $220\sigma$ ( $180\sigma$ ) |
| Number of SZ clusters  | 20000                       | 33000                       |
| Kinematic SZ detection (DESI cross-correlation)                        | $190\sigma$                 | $240\sigma$                 |
| Measurement of Optical Depth from kSZ, $\sigma(\tau)$ <sup>b</sup>     | 0.007                       | 0.003                       |
| <b>Primordial polarization (SATs)</b>                                  |                             |                             |
| Tensor-to-scalar ratio   | $\sigma(r) = 0.002$         | $\sigma(r) = 0.001^c$       |

# SO Pipeline Status

- Validated analysis pipeline for B-mode search based on SAT instrumentation data



[SO large scale B-mode pipeline validation paper](#)

- LAT data analysis pipelines have been tested on recent ACT data release (DR6)